



Evidence for Sulfur-bearing Species on Callisto's Leading Hemisphere: Sourced from Jupiter's Irregular Satellites or Io?

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Abstract

We investigated whether sulfur-bearing species are present on the icy Galilean moon Callisto by analyzing eight near-infrared reflectance spectra collected over a wide range of sub-observer longitudes. We measured the band areas and depths of a $4\ \mu\text{m}$ feature in these spectra, which has been attributed to sulfur dioxide (SO_2), as well as carbonates, in previously collected data sets of this moon. All eight spectra we collected display the $4\ \mu\text{m}$ band. The four spectra collected over Callisto's leading hemisphere display significantly stronger $4\ \mu\text{m}$ bands compared to the four trailing hemisphere spectra ($>3\sigma$ difference). We compared the central wavelength position and shape of Callisto's $4\ \mu\text{m}$ band to laboratory spectra of various sulfur-bearing species and carbonates. Our comparison demonstrates that Callisto's $4\ \mu\text{m}$ band has a spectral signature similar to thermally altered sulfur, as well as a $4.025\ \mu\text{m}$ feature attributed to disulfanide (HS_2). Our analysis therefore supports the presence of S-bearing species on Callisto but is not consistent with the presence of SO_2 . The significantly stronger $4\ \mu\text{m}$ band detected on Callisto's leading hemisphere could result from collisions with H_2S -rich dust grains that originate on Jupiter's retrograde irregular satellites or implantation of magnetospheric S ions that originate from volcanic activity on Io. Alternatively, S-bearing species could be native to Callisto and are exposed by dust collisions and larger impacts that drive regolith overturn, primarily on its leading side.

Unified Astronomy Thesaurus concepts: Planetary surfaces (2113); Surface composition (2115); Surface processes (2116); Surface ices (2117)

1. Introduction

Sulfur is one of the key elements required for life as we know it. Measuring the abundance, distribution, and spectral signature of sulfur on ocean worlds like the icy Galilean moons Callisto, Ganymede, and Europa is therefore important for assessing their astrobiological potential (e.g., Hendrix et al. 2019). The innermost Galilean moon Io erupts substantial quantities of S-bearing neutrals into orbit, which are ionized and subsequently trapped in Jupiter's magnetosphere (e.g., Schneider & Bagenal 2007). These sulfur ions are delivered primarily to the trailing hemispheres of Europa and Ganymede by Jupiter's corotating plasma, spurring a cascade of radiolytic surface chemistry (e.g., Johnson et al. 2004). The flux of magnetospheric S ions is lower at the orbit of Callisto but is still sufficient to drive radiolytic modification of its surface.

Analysis of spectra collected by the Near Infrared Mapping Spectrometer (NIMS) on the Galileo spacecraft has revealed the presence of an absorption band centered near $4.25\ \mu\text{m}$, which likely results from carbon dioxide (CO_2 ; e.g., Carlson et al. 1996; McCord et al. 1997, 1998). This $4.25\ \mu\text{m}$ CO_2 band is distributed across Callisto's trailing hemisphere, producing a bullseye shaped pattern centered around 270° longitude (i.e., the center of its trailing hemisphere; e.g., Hibbitts et al. 2000) and is likely formed by Jupiter's corotating plasma flowing onto Callisto's trailing side. CO_2 is also present on Callisto's

leading hemisphere, but displays a different distribution, with concentrated deposits of CO_2 spatially associated with impact craters and their ejecta blankets (Hibbitts et al. 2000, 2002).

Other prominent absorption bands are present in NIMS spectra of Callisto, including a feature centered between 4.02 and $4.05\ \mu\text{m}$ (hereafter referred to as the “ $4\ \mu\text{m}$ ” band), which was originally attributed to sulfur dioxide (SO_2) on Europa and Ganymede, as well as on Callisto (e.g., McCord et al. 1997, 1998; Hibbitts et al. 2000). Analysis of NIMS spectra demonstrated that the $4\ \mu\text{m}$ band is strongest on the trailing hemispheres of Europa and Ganymede, consistent with implantation of S ions from Io. In contrast, Callisto's $4\ \mu\text{m}$ band appears to be stronger on its leading hemisphere, hinting at a possibly different origin from the $4\ \mu\text{m}$ band detected on Europa and Ganymede (Hibbitts et al. 2000). The $4\ \mu\text{m}$ band has a mottled distribution across Callisto's leading hemisphere, without clear association with impact features or other landforms, unlike CO_2 (e.g., Hibbitts et al. 2000), suggesting that the $4\ \mu\text{m}$ feature could have a different origin from CO_2 on this moon.

In contrast to the sharp shape and strong absorption exhibited by the $4.07\ \mu\text{m}$ SO_2 frost combination band ($\nu_1 + \nu_3$) (e.g., Nash & Betts 1995) detected on Io (e.g., Fanale et al. 1979; Cruikshank 1980), Callisto's $4\ \mu\text{m}$ band is more subtle, has a rounded shape, and is shifted to shorter wavelengths ($\sim 4.02\text{--}4.05\ \mu\text{m}$). These spectral differences have led some studies to cast doubt on the presence of S-bearing species like SO_2 on Callisto, suggesting instead that C-rich species like calcium carbonate (CaCO_3) or sodium carbonate (Na_2CO_3)

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might be better candidates to explain the $4\ \mu\text{m}$ band (Johnson et al. 2004). Similarly, analysis of data collected by the Ultraviolet Spectrometer (UVS) on board Galileo indicates that the $0.28\ \mu\text{m}$ feature detected on Callisto’s leading side is more consistent with the spectral slope of C-rich material (Hendrix & Johnson 2008), as opposed to an SO_2 band suggested by prior work (e.g., Lane & Domingue 1997; Noll et al. 1997).

Therefore, the origin and composition of Callisto’s $4\ \mu\text{m}$ band, and the possible presence of sulfur on this moon, remain uncertain. Complicating matters, the often low signal-to-noise (S/N) and low resolving power ($R \sim 200$) of NIMS spectra has made assessment of the spectral signature and distribution of Callisto’s $4\ \mu\text{m}$ band more challenging. To investigate the $4\ \mu\text{m}$ band further, we collected new near-infrared (NIR) reflectance spectra of Callisto. We detected the $4\ \mu\text{m}$ band in these new ground-based spectra, along with other absorption bands between 2.8 and $5.0\ \mu\text{m}$. We used these data to characterize the spectral signature and longitudinal distribution of Callisto’s $4\ \mu\text{m}$ band, providing new constraints on the presence and origin of S-bearing species.

2. Data and Methods

2.1. Observations and Data Reduction

We observed Callisto on eight different nights in 2020 May and June using the NIR SpeX spectrograph/imager (e.g., Rayner et al. 2003) on NASA’s Infrared Telescope Facility (IRTF), operating in long cross-dispersed mode, spanning 1.98 – $5.3\ \mu\text{m}$ over seven spectral orders (“LXD_long” mode). All LXD_long observations were made with a $0''.3$ wide by $1''$ long slit, providing an average resolving power of 2500 in each spectral order. Prior to each observation, the SpeX slit was oriented parallel to Callisto’s poles, and then placed over the center of Callisto’s disk, thereby ensuring that the sub-observer longitude was fully covered by the slit. These spectra were collected using “AB” nodding, where all targets are observed in two different positions, separated by $7''.5$. Each “A” and “B” frame had a maximum exposure length of 3 s to prevent saturation at wavelengths $>4.1\ \mu\text{m}$. During data reduction, these “A” and “B” frames were separated into sequential pairs, and the “B” frames were subtracted from the “A” frames, thereby performing first-order sky emission correction.

Flat-field frames were generated with an internal quartz lamp to correct for pixel variations across the chip. To perform wavelength calibration of these data, we collected arc lamp frames, using an internal argon lamp. Because argon emission lines are weak in L/L’ and M bands (~ 2.8 – $5.3\ \mu\text{m}$), we utilized night sky emission lines to calibrate the longer wavelength data. Data extraction, flat-fielding, wavelength calibration, and spectral order merging were conducted using the Spextool data reduction suite (Cushing et al. 2004), as well as custom programs. All extracted spectra were divided by standard star spectra, observed on the same night immediately before and after observations of Callisto (within ± 0.1 airmass), to provide additional telluric correction and remove the solar spectrum. The standard stars we observed were HD 189893 (G0V) and SA 112 1333 (F8). After dividing each spectrum by the standard star, all Callisto frames collected on the same night were coadded to boost S/N. The observation details are summarized in Table 1.

2.2. Band Parameter Analyses

To assess the spectral signature and longitudinal distribution of the $4\ \mu\text{m}$ band, we measured its center, area, and depth in each collected spectrum with a custom band parameter analysis program that our team has used previously to analyze reflectance data of icy moons (e.g., Cartwright et al. 2015, 2020b, 2020a). After reading in the data, the program fit the continuum of each spectrum between 3.926 and $4.114\ \mu\text{m}$ to measure the $4\ \mu\text{m}$ band. Next, the program divided each $4\ \mu\text{m}$ band by its continuum, and measured the area of the resulting continuum-divided band using the trapezoidal rule. The program then used Monte Carlo simulations to estimate the uncertainties for these band area measurements by resampling the 1σ errors of each spectral channel within each band (iterated 20,000 times). To measure the depth of each spectrum’s $4\ \mu\text{m}$ band, the program identified the spectral channel closest to $4.02\ \mu\text{m}$ as the band center, and then calculated the mean reflectance using all spectral channels within $\pm 0.015\ \mu\text{m}$. To calculate the band depth, the program subtracted these mean reflectances from 1. To estimate measurement errors, the program added the 1σ uncertainties of all spectral channels included in the mean reflectance measurement, in quadrature, and then divided by the number of channels (n) to calculate the mean uncertainty ($\bar{\sigma}$). To estimate the point-to-point variation in each band, the program calculated the standard deviation of the mean ($\sigma_{\bar{x}} = \sigma/\sqrt{n}$). The program then calculated the final uncertainty for each band depth measurement by adding $\bar{\sigma}$ and $\sigma_{\bar{x}}$ in quadrature. An example of our band measurement procedure is shown in Figure A1.

3. Results and Analysis

3.1. IRTF/SpeX Spectra of Callisto

All eight of the SpeX spectra we collected display evidence for the $4\ \mu\text{m}$ band, centered near $4.02\ \mu\text{m}$ (Figures 1, A2). Our $4\ \mu\text{m}$ band area and depth measurements are reported in Table 2 (along with their 1σ uncertainties). These measurements confirm the presence of the $4\ \mu\text{m}$ band in all eight spectra, demonstrating that this feature is present across Callisto’s surface. Furthermore, the four spectra collected over Callisto’s leading hemisphere (1° – 180° longitude) show significantly stronger $4\ \mu\text{m}$ bands compared to the four spectra collected over its trailing hemisphere (181° – 360° longitude), which is consistent with previous analyses (e.g., Hibbitts et al. 2000; see our Figure A3). We further investigate the spatial distribution of the $4\ \mu\text{m}$ band in Section 3.2.

The SpeX spectra we collected display other spectral features that we briefly describe here. All collected SpeX spectra display the strong $3\ \mu\text{m}$ H_2O ice band, as well as a subtle $3.1\ \mu\text{m}$ Fresnel peak (e.g., Mastrapa et al. 2009). These H_2O ice features were detected previously in ground-based (e.g., Pollack et al. 1978; Calvin & Clark 1993) and NIMS (e.g., Carlson et al. 1996) spectra. The SpeX data display a wide band centered near $4.57\ \mu\text{m}$, hereafter referred to as the “ $4.6\ \mu\text{m}$ ” band, which was originally detected by NIMS and attributed to CN-bearing species (e.g., McCord et al. 1998), or alternatively, carbon suboxide (C_3O_2 , Johnson et al. 2004). Similar to the $4\ \mu\text{m}$ band, the SpeX spectra we collected indicate that the $4.6\ \mu\text{m}$ band is stronger on Callisto’s leading hemisphere. The SpeX spectra also display a band centered near $3.87\ \mu\text{m}$, hereafter referred to as the “ $3.9\ \mu\text{m}$ ” band. This band was previously detected by NIMS and attributed to SH-bearing materials (e.g., McCord et al. 1998), as well as

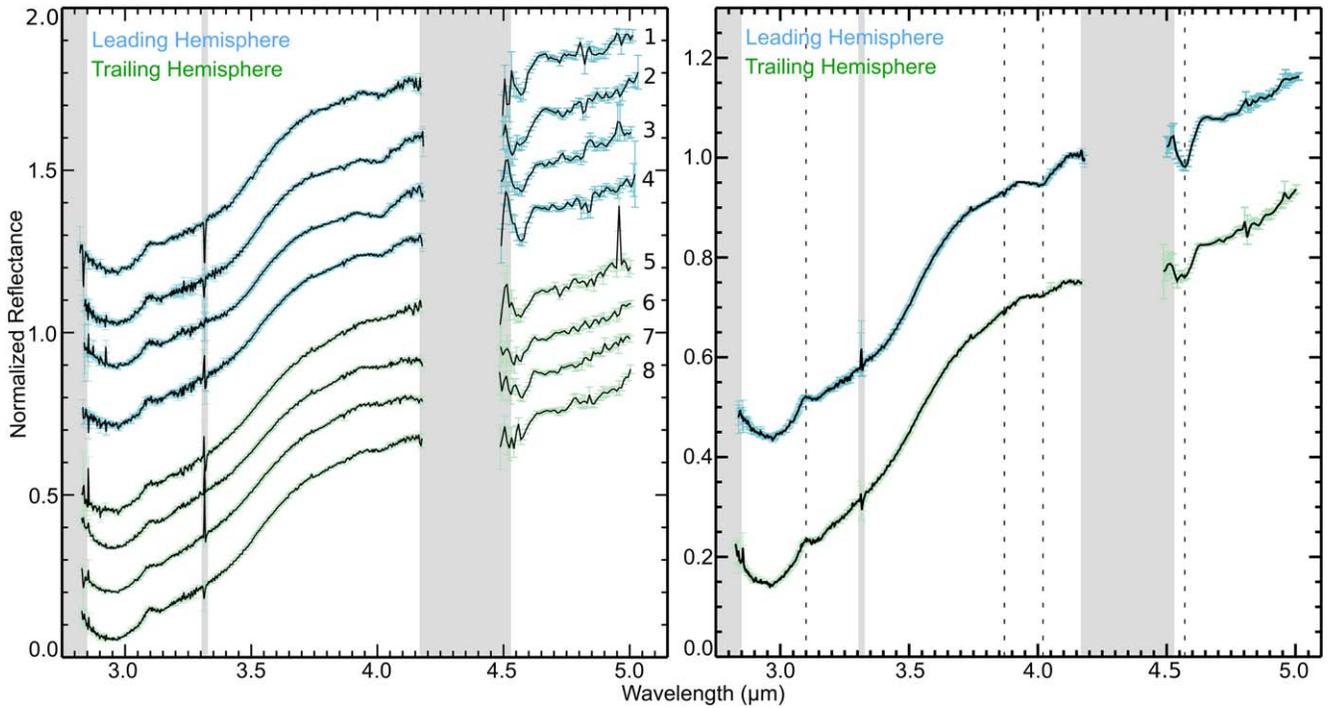


Figure 1. Left: eight SpeX spectra of Callisto’s leading and trailing hemisphere (1σ uncertainties shown as blue and green error bars, respectively), normalized to 1 at $4.12\ \mu\text{m}$ and offset vertically for clarity. The spectra are numbered (1–8) using the same sequence shown in Tables 1 and 2. Right: “grand average” leading and trailing hemisphere spectra of Callisto (1σ uncertainties shown as blue and green error bars, respectively), normalized to 1 at $4.12\ \mu\text{m}$ and offset vertically for clarity. The dashed lines at 3.1 , 3.87 , 4.02 , and $4.57\ \mu\text{m}$ highlight the central wavelengths of the H_2O ice Fresnel peak and the $3.9\ \mu\text{m}$, $4\ \mu\text{m}$, and $4.6\ \mu\text{m}$ bands, respectively. The $4\ \mu\text{m}$ and $4.6\ \mu\text{m}$ bands are stronger on Callisto’s leading hemisphere, whereas the $3.9\ \mu\text{m}$ band does not display apparent hemispherical asymmetries in its band strength. Other, more subtle bands that are possibly present include features centered near 2.97 , 3.05 , and $3.4\ \mu\text{m}$, which appear to be stronger on Callisto’s leading hemisphere, and a band centered near $3.75\ \mu\text{m}$, which appears to be slightly stronger on Callisto’s trailing side (see Figure A4 for closer looks at these subtle bands). All spectra have been lightly smoothed using a 9 pixel wide boxcar function between 2.75 and $4.2\ \mu\text{m}$ and a 27 pixel wide boxcar function between 4.45 and $5.05\ \mu\text{m}$. Wavelength regions where strong telluric bands are present are shown as gray-toned zones.

Table 1
IRTF/SpeX Observations of Callisto

Spectrum Number	Sub-observer Longitude ($^\circ$)	Sub-observer Latitude ($^\circ$)	Angular Diameter ($''$)	Phase Angle ($^\circ$)	UT Date	UT Time (mid-expos)	Integration Time (minutes)	Airmass Range
1	20.7	−1.0	1.56	5.78	20/06/15	13:00	24	1.33–1.39
2	64.0	−1.0	1.56	5.45	20/06/17	13:00	25	1.33–1.46
3	120.6	−1.1	1.52	7.75	20/06/03	12:45	26	1.32–1.49
4	164.7	−1.1	1.53	7.44	20/06/05	13:30	25	1.31–1.39
5	200.7	−1.1	1.47	9.37	20/05/21	13:45	24	1.32–1.40
6	244.0	−1.1	1.47	9.13	20/05/23	13:45	25	1.32–1.39
7	287.2	−1.1	1.48	8.88	20/05/25	13:45	28	1.33–1.41
8	337.8	−1.0	1.55	6.11	20/06/13	13:15	26	1.33–1.42

carbonic acid (H_2CO_3 , e.g., Hage et al. 1998). The $3.9\ \mu\text{m}$ band in our SpeX spectra does not appear to display strong leading/trailing asymmetries in its distribution, and it is notably weaker than the $3.9\ \mu\text{m}$ band detected by NIMS. Hints of other, more subtle bands are present in these spectra as well, which we discuss in Appendix A.4 (Figure A4). Although of great interest, quantitative analysis of these other absorption features is beyond the scope of this paper and will be included in future work.

3.2. Longitudinal Distribution of the $4\ \mu\text{m}$ Band

Previous work has demonstrated that Callisto displays longitudinal trends in its surface composition, with weaker

H_2O ice bands and more red material on its leading hemisphere (e.g., Morrison et al. 1974), and more CO_2 on its trailing hemisphere (e.g., Hibbitts et al. 2000). Consequently, we searched for longitudinal trends in the distribution of the $4\ \mu\text{m}$ band on Callisto by calculating the mean $4\ \mu\text{m}$ band areas for Callisto’s leading and trailing hemispheres and comparing them: $28.35 \pm 4.37\ \mu\text{m}^{-4}$ and $4.20 \pm 0.37\ \mu\text{m}^{-4}$, respectively. We also calculated and compared the mean $4\ \mu\text{m}$ band depths for Callisto’s leading and trailing hemispheres: $0.029 \pm 0.008\ \mu\text{m}$ and $0.007 \pm 0.001\ \mu\text{m}$, respectively. Comparison of these values demonstrates that the mean $4\ \mu\text{m}$ band is significantly stronger on Callisto’s leading hemisphere compared to its trailing hemisphere ($>4\sigma$ difference).

Table 2
4 μm Band Parameter Measurements

Spectrum Number	Sub-observer Longitude ($^{\circ}$)	Sub-observer Latitude ($^{\circ}$)	Longitudinal Quadrant (Geologic Region)	Band Area ($10^{-4} \mu\text{m}$)	Band Depth
1	20.7	-1.0	Sub-Jupiter (Valhalla)	24.0 ± 0.7	0.026 ± 0.002
2	64.0	-1.0	Leading (Valhalla)	33.4 ± 0.5	0.033 ± 0.001
3	120.6	-1.1	Leading (Asgard)	37.6 ± 0.5	0.038 ± 0.001
4	164.7	-1.1	Anti-Jupiter (Asgard)	18.4 ± 0.5	0.021 ± 0.001
5	200.7	-1.1	Anti-Jupiter	4.7 ± 0.5	0.006 ± 0.001
6	244.0	-1.1	Trailing	4.2 ± 0.4	0.007 ± 0.001
7	287.2	-1.1	Trailing	3.7 ± 0.4	0.008 ± 0.001
8	337.8	-1.0	Sub-Jupiter	9.4 ± 0.5	0.012 ± 0.001

To further investigate the distribution of the 4 μm band, we plotted the eight individual band area and depth measurements as a function of longitude (Figure 2). We fit the band area and depth measurements with a sinusoidal model, representing a surface displaying longitudinal trends in the distribution of the 4 μm band, and a mean model, representing a surface without leading/trailing trends in the distribution of this feature. We used an F -test (e.g., Lomax & Hahs-Vaughn 2013) to compare these two models. Our null hypothesis was that there is no difference between these two model fits, indicating that there is no meaningful longitudinal difference in the distribution of the 4 μm band. Our F -test analysis demonstrates that there is a $>3\sigma$ difference between the mean and sinusoidal models for both the 4 μm band area and depth measurements ($p < 0.002$, Table A1).

Thus, the mean 4 μm band areas and depths, as well as the results of our F -test analysis, demonstrates that the 4 μm band is significantly stronger on Callisto's leading hemisphere. The 4 μm band is strongest in Spectrum 3 (mid-observation, sub-observer longitude 120.6°), which is located between the center of the Callisto's leading hemisphere (longitude 90°) and the Asgard impact basin (centered near 140° longitude). Conversely, the 4 μm band is weakest in Spectrum 7 (mid-observation, sub-observer longitude of 287.2°), which is proximal to the center of Callisto's trailing side (longitude 270°). We consider the implications of the longitudinal distribution of the 4 μm band in Section 4.

3.3. Comparing Callisto Spectra to Laboratory Spectra of Candidate Species

To investigate the constituents contributing to the 4 μm band, we compared Callisto's grand average leading and trailing hemisphere spectra to laboratory reflectance spectra of candidate species (Figure 3). We first focused on sulfur-bearing materials that have been previously suggested (summarized in Johnson et al. 2004; Moore et al. 2004), including SO_2 mixed with H_2O ice (1:1 ratio; Moore et al. 2002), SO_2 adsorbed on cabosil (Nash & Betts 1995), and thermally altered sulfur (Fanale et al. 1979). We also compared Callisto's 4 μm band to a spectral feature identified near $4.025 \mu\text{m}$, formed by irradiation of substrates composed of H_2O ice and hydrogen sulfide (H_2S ; Salama et al. 1990; Moore et al. 2007; Jiménez-Escobar & Muñoz Caro 2011), which has been attributed to disulfanide (HS_2 ; Jiménez-Escobar & Muñoz Caro 2011). The $4.025 \mu\text{m}$ feature was extracted from a transmission (T) spectrum calculated from absorbance (a): $T = 10^{-a}$, measured at a temperature of 133 K (Jiménez-Escobar & Muñoz Caro 2011). Additionally, we compared Callisto's 4 μm band

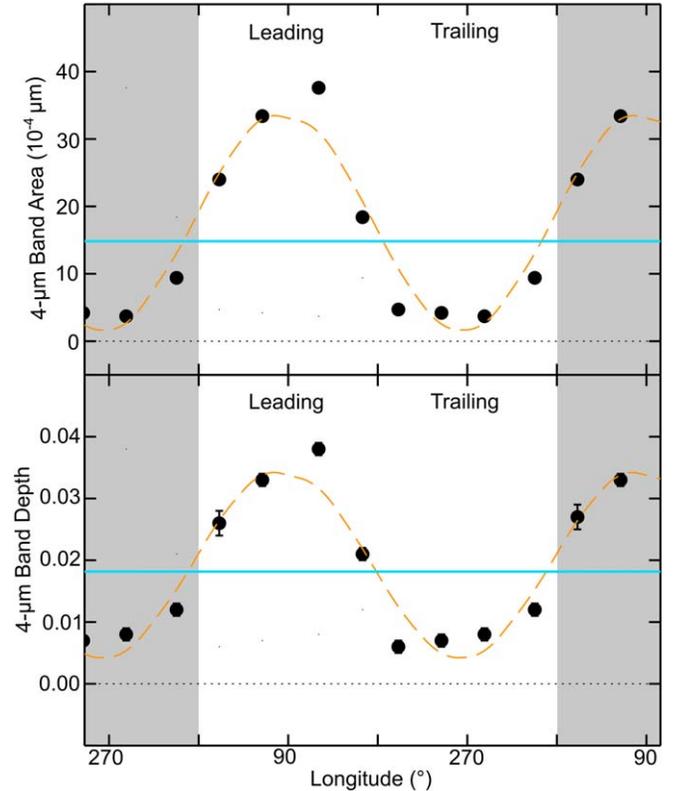


Figure 2. 4 μm band area (top plot) and band depth (bottom plot) measurements and 1σ uncertainties for all eight Callisto spectra, shown as a function of sub-observer longitude (Table 2). For the band area measurements, and some of the band depth measurements, the 1σ uncertainties are smaller than the size of the black-filled circles. We fit these data with a sinusoidal model (orange dashed line), which represents a surface with significant variations in the abundance of this constituent, and the mean 4 μm band measurement (blue line), representing a surface that displays no variation in the abundance of this constituent. The maxima of the sinusoidal models are not locked to specific longitudes. Duplicate longitudes are shown to highlight the periodicity in the distribution of this constituent (gray-toned zones). We compared the model fits using an F -test, finding that the 4 μm band displays significant longitudinal variations in its distribution ($p < 0.002$, F -test results summarized in Appendix A.5).

to laboratory reflectance spectra of the carbonates Na_2CO_3 and CaCO_3 (Nyquist et al. 1997), which have prominent 4 μm bands (Johnson et al. 2004). Previous studies have also considered carbonic acid, methanol, and other species, finding that these constituents do not provide good matches to the 4 μm band (e.g., Hudson & Moore 2001; Hand et al. 2007).

Our comparison demonstrates that SO_2 -bearing materials are poor matches to Callisto's 4 μm band, both in terms of band center and shape (Figure 3). In contrast, thermally altered sulfur

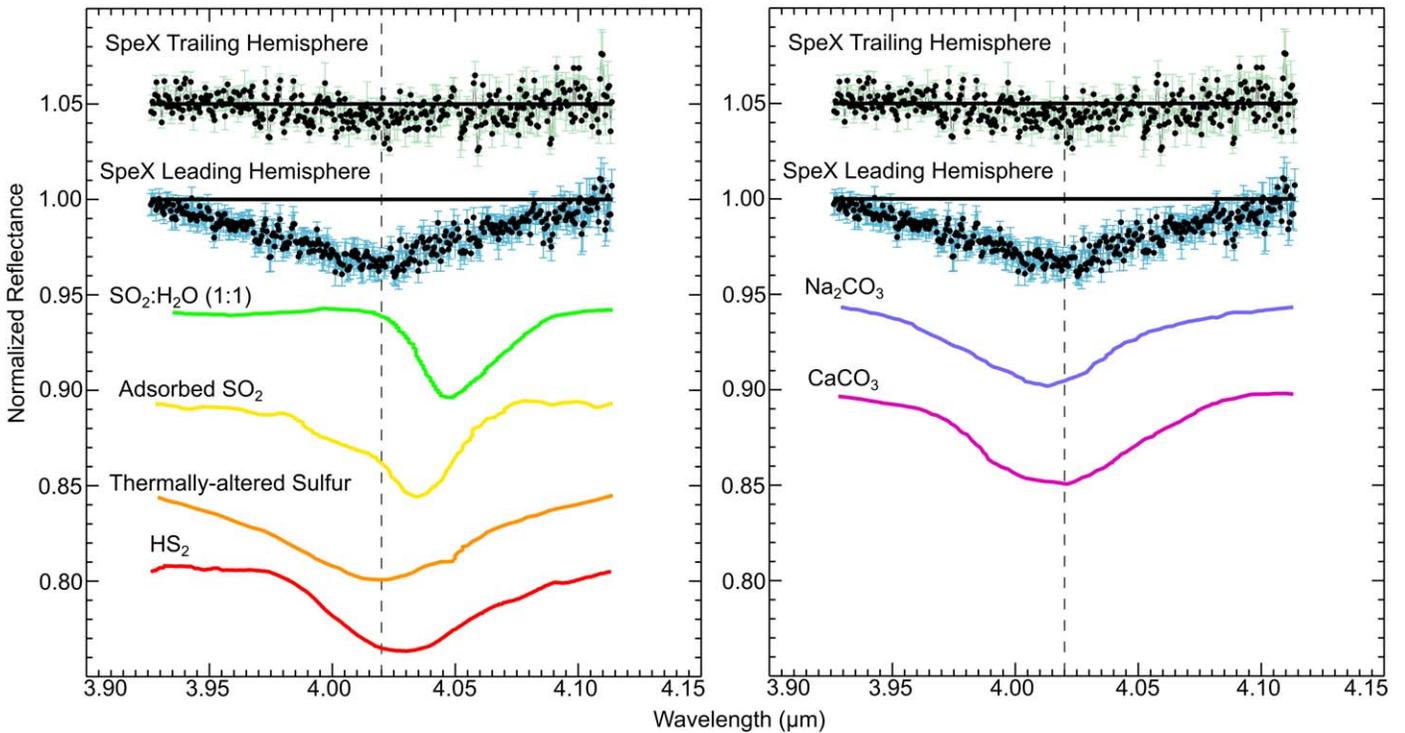


Figure 3. Left: comparison between continuum-divided $4 \mu\text{m}$ bands identified in the grand average leading and trailing hemisphere spectra of Callisto (shown in Figure 1) and continuum-divided laboratory spectra of SO_2 in H_2O ice (green; Moore et al. 2002), adsorbed SO_2 (gold; Nash & Betts 1995), thermally altered S (orange; Fanale et al. 1979), and a transmission feature centered near $4.025 \mu\text{m}$ (red) that has been attributed to HS_2 (Jiménez-Escobar & Muñoz Caro 2011), offset vertically for clarity. Right: comparison between Callisto’s continuum-divided $4 \mu\text{m}$ bands and continuum-divided laboratory spectra of Na_2CO_3 and CaCO_3 (violet and purple, respectively, Nyquist et al. 1997), offset vertically for clarity. In both plots, the dashed line at $4.02 \mu\text{m}$ marks the central wavelength position of the $4 \mu\text{m}$ band identified in the SpeX spectra of Callisto.

provides a reasonable match in terms of band center and shape. The spectrum shown in Figure 3 is from a sample heated to $\sim 473 \text{ K}$ to produce polymeric sulfur, which likely includes S_2 , S_3 , or other S allotropes (Fanale et al. 1979). Although this laboratory sample was formed at substantially higher temperatures than Callisto’s estimated peak surface temperature ($\sim 165 \text{ K}$; Spencer 1987), S allotropes could also be formed by radiolytic modification of Callisto surface. For example, Iogenic S ions delivered to Callisto in Jupiter’s magnetosphere should react with its surface constituents, spurring a radiolytic production cycle that could generate S allotropes, similar to the sulfur cycle operating on Europa (e.g., Carlson et al. 2002).

The $4.025 \mu\text{m}$ feature, attributed to HS_2 , also provides a reasonable match to the band center and shape of Callisto’s $4 \mu\text{m}$ band. This $4.025 \mu\text{m}$ feature was formed in the laboratory by irradiating mixtures of H_2O ice and H_2S at cryogenic temperatures ($< 90 \text{ K}$). After irradiation, these samples were warmed up to 224 K , and the $4.025 \mu\text{m}$ feature persisted over the entire temperature range, whereas the H_2S frost fundamental band centered near $3.9 \mu\text{m}$ (ν_1) started to decrease once the sample chamber reached temperatures $> 100 \text{ K}$ (see Figure 6 in Jiménez-Escobar & Muñoz Caro 2011). Therefore, at Callisto’s peak surface temperature ($\sim 165 \text{ K}$), H_2S should sublime, thereby weakening the $3.9 \mu\text{m}$ H_2S fundamental band, whereas the $4.025 \mu\text{m}$ feature attributed to HS_2 should persist, possibly explaining the weakness of the $3.9 \mu\text{m}$ band and stronger $4 \mu\text{m}$ band we detected with SpeX. Of note, UV irradiation of H_2S ice should also lead to the formation of S allotropes like S_2 and S_3 (Jiménez-Escobar & Muñoz Caro 2011).

Carbonates like Na_2CO_3 and CaCO_3 also provide good matches to the $4 \mu\text{m}$ feature in terms of central wavelength position and band shape, as previously demonstrated (Johnson et al. 2004). However, carbonates display prominent absorption features at shorter wavelengths (between $\sim 2.2\text{--}2.4$ and $\sim 3.3\text{--}3.5 \mu\text{m}$), which we do not detect in the SpeX spectra of Callisto (Figures 1 and A2) and were not detected by NIMS, reducing the likelihood that carbonates are the primary contributors to the $4 \mu\text{m}$ band. Nevertheless, the surface of Callisto has a vigorous carbon system, and perhaps C- and S-bearing species, along with H_2O ice, participate in a shared radiolytic cycle, similar to sulfur and H_2O ice on Europa (e.g., Carlson et al. 2002). Such a C–S– H_2O radiolytic cycle might produce species like carbonyl sulfide (OCS) and carbon disulfide (CS_2), which display absorption features at 3.4 and $4.6 \mu\text{m}$, respectively (Coblentz Society 1964), possibly explaining the presence of these two features in NIMS and SpeX spectra of Callisto.

In summary, we find that thermally altered sulfur provides a good match to the spectral signature of Callisto’s $4 \mu\text{m}$ band. The $4.025 \mu\text{m}$ feature attributed to HS_2 can also explain the presence of the $4 \mu\text{m}$ band. Thus, our analysis supports the presence of S-bearing constituents on Callisto, but we do not find strong evidence for the previously suggested SO_2 .

4. Discussion and Conclusions

We analyzed eight SpeX spectra of Callisto, gaining new knowledge on the spectral signature and distribution of the $4 \mu\text{m}$ band. We identified a leading/trailing hemispherical trend in the distribution of the $4 \mu\text{m}$ band, with significantly

stronger band measurements on Callisto’s leading side. Our analysis suggests that the spectral signature of Callisto’s $4\ \mu\text{m}$ band is consistent with thermally altered sulfur and a feature centered near $4.025\ \mu\text{m}$, which has been attributed to HS_2 (Jiménez-Escobar & Muñoz Caro 2011). Below, we consider possible origin scenarios to explain the presence of the $4\ \mu\text{m}$ band on Callisto.

Delivery of S-bearing dust from the irregular satellites: Dust grains ($\sim 10\text{--}1000\ \mu\text{m}$ diameters) ejected from the giant planets’ retrograde irregular satellites should experience Poynting–Robertson drag and slowly migrate inward on decaying orbits (e.g., Burns et al. 1979). Eventually, the orbits of these dust grains overlap the orbital zone of the classical satellites, and they subsequently collide with the leading hemispheres of the outermost moons (e.g., Bottke et al. 2013; Tamayo et al. 2013). The irregular satellites of the giant planets are darker and spectrally redder than the classical moons (e.g., Graykowski & Jewitt 2018). The accumulation of irregular satellite dust can therefore explain why the leading hemispheres of the outer classical moons Callisto (e.g., Pollack et al. 1978), Iapetus (e.g., Cruikshank et al. 1983), and Titania and Oberon (e.g., Cartwright et al. 2018) are spectrally redder and darker than their trailing hemispheres.

Similar to the irregular satellites, Jupiter’s trojan asteroid population includes a group of objects with reddish surfaces (e.g., Szabó et al. 2007; Emery et al. 2010). It has been hypothesized that this red-colored trojan asteroid group is populated by captured Kuiper Belt objects (KBOs) that formed beyond an H_2S snow line in the primordial Kuiper Belt (Wong & Brown 2016). Dynamical simulations indicate that giant planet migration in the early solar system scattered large numbers of KBOs into the giant planet zone, and some of these objects were likely captured by Jupiter into its trojan asteroid and irregular satellite populations (e.g., Morbidelli & Nesvorný 2020). Along with capture of H_2S -rich KBOs in Jupiter’s L4 and L5 Lagrange points, perhaps H_2S -bearing KBOs were captured into Jupiter’s irregular satellite population, providing a source of H_2S and other S-bearing species that could be delivered to Callisto’s leading hemisphere within dust grains. Although sublimation and space weathering should remove volatile-rich deposits from the surfaces of Jovian irregular satellites, H_2S and other volatiles could be retained beneath their regoliths and subsequently excavated and ejected in dust grains by impact events. Albeit, it is uncertain how long H_2S could survive within dust grains before colliding with Callisto, nor whether H_2S delivered to Callisto would survive long enough to be radiolytically modified into other S-bearing species. Laboratory work that explores the longevity of H_2S under conditions relevant to the surface of Callisto are needed to further investigate this scenario.

Delivery of magnetospheric S ions: Volcanoes on Io erupt large volumes of S-rich material into orbit. Much of the erupted

sulfur is then ionized and gets trapped in Jupiter’s magnetosphere (e.g., Schneider & Bagenal 2007). These trapped S ions are delivered to the more distant icy Galilean moons, primarily bombarding the trailing hemispheres of Europa and Ganymede (e.g., Johnson et al. 2004). In contrast, recent numerical modeling work shows that energetic O and S ions ($1\text{--}100\ \text{KeV}$) trapped in Jupiter’s magnetosphere might preferentially bombard Callisto’s leading hemisphere (Liuzzo et al. 2019). This hemispherical dichotomy results from interactions between Jupiter’s magnetosphere and Callisto’s intense ionosphere (Kliore et al. 2002), which largely blocks these lower energy O and S ions from accessing its trailing side, unlike Europa and Ganymede, which do not have substantial ionospheres. Radiolytic sulfur chemistry, induced by implanted S ions, could therefore be operating across Callisto’s leading hemisphere, explaining the stronger $4\ \mu\text{m}$ band we detected on this hemisphere.

Exposure of native S-rich deposits: Alternatively, the species contributing to the $4\ \mu\text{m}$ band could be native to Callisto and are exposed by dust-driven regolith overturn, which should preferentially operate on the leading hemispheres of tidally locked moons (e.g., Bennett et al. 2013). In this scenario, native deposits of S-bearing species exposed by dust collisions could be modified by UV photolysis and charged particle radiolysis, forming new S-rich constituents, such as S_2 -bearing species and perhaps CS-bearing species as well. Larger impact events might also excavate S-rich deposits from greater depths in Callisto’s subsurface, albeit, analysis of NIMS data indicates that the $4\ \mu\text{m}$ band is not strongly spatially associated with craters on Callisto’s leading hemisphere, unlike CO_2 (Hibbitts et al. 2000).

To further test these different origin hypotheses, follow-up ground-based observations of the $4\ \mu\text{m}$ band at complementary sub-observer longitudes are needed. Furthermore, new laboratory experiments that measure the spectral signature of substrates composed of H_2O ice, S-bearing, and C-bearing species, performed under temperature and irradiation conditions relevant to the surface of Callisto, could provide key insight into origin and composition of the $4\ \mu\text{m}$ band. These new experiments could also provide insight into whether a radiolytic cycle involving S, C, and H_2O ice is occurring on Callisto, spurring the formation of species like OCS and CS_2 .

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Appendix

A.1. Methods: Band Parameter Measurements

In this appendix, we show an example of our band parameter measurement technique (Figure A1).

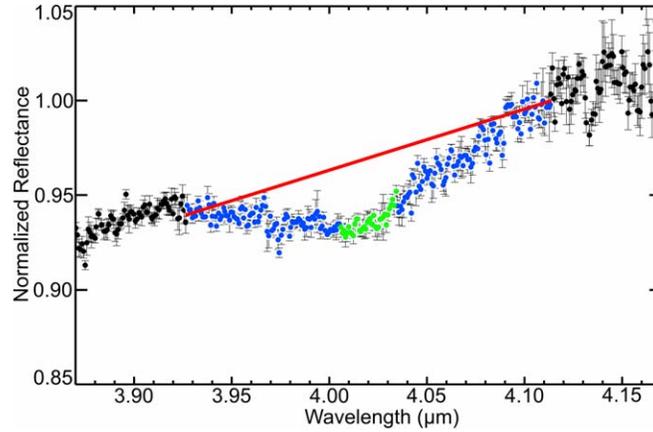


Figure A1. An example of our $4\ \mu\text{m}$ band area and depth measurement procedure, using a SpeX spectrum of Callisto collected at mid-observation, sub-observer longitude 120.6° (Spectrum 3 in Figure 1 and Tables 1 and 2). The data points used for our band area measurement (blue, $3.926\text{--}4.114\ \mu\text{m}$) and those used for our band depth measurement (green, $4.005\text{--}4.035\ \mu\text{m}$) are highlighted. The $4\ \mu\text{m}$ band continuum was calculated by averaging the data points within $\pm 0.01\ \mu\text{m}$ of 3.926 and $4.114\ \mu\text{m}$, and we then connected these two mean continuum reflectances with a line (red).

A.2. Results: IRTF/SpeX Spectra

In this appendix, we present eight near-infrared reflectance spectra of the Galilean moon Callisto at their native spectral resolution (Figure A2).

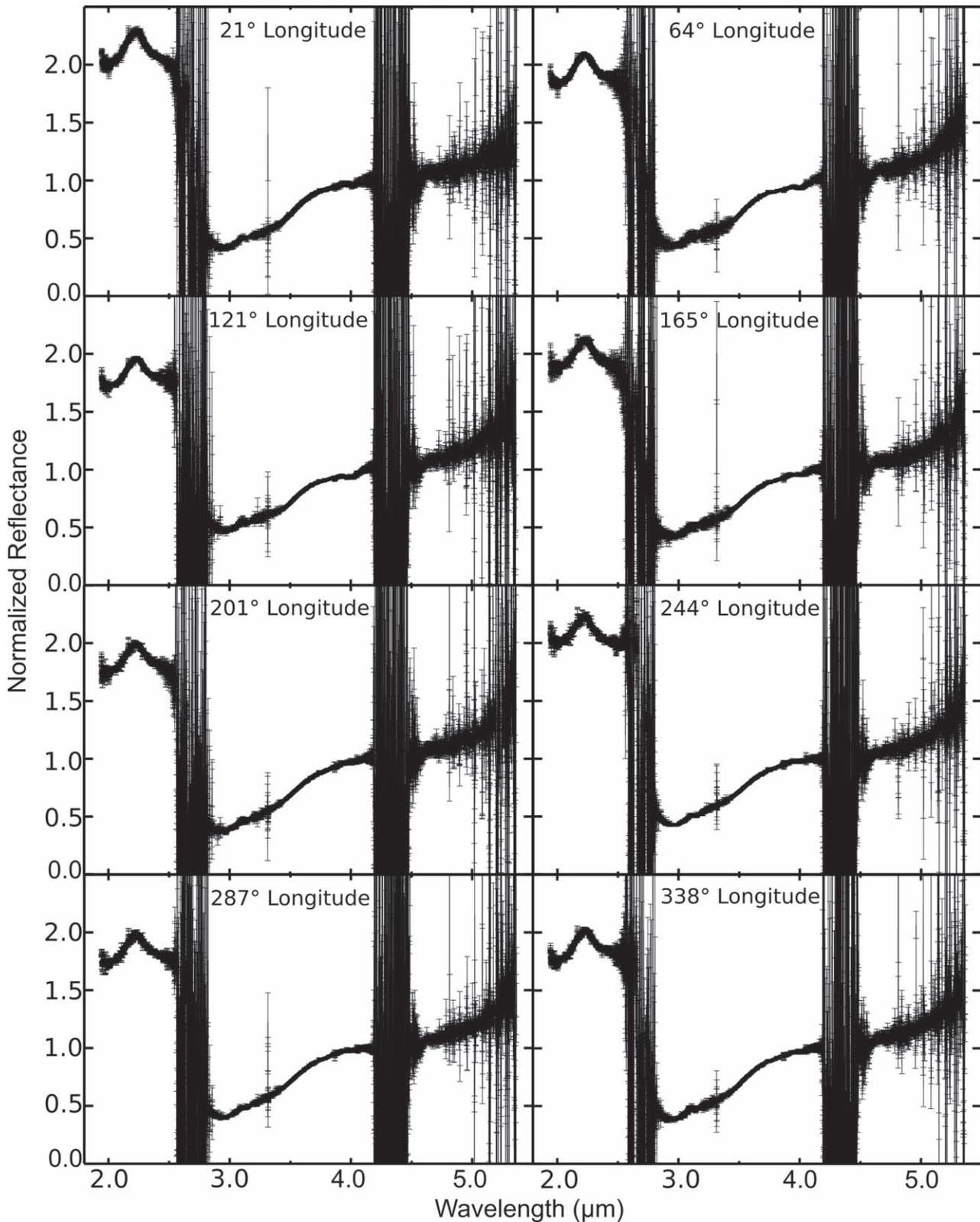


Figure A2. Eight new SpeX spectra of Callisto and their 1σ uncertainties, collected in 2020 May and June. The mid-observation, sub-observer longitude for each spectrum is listed along the top of each plot (see Table 1 for observation details). All spectra have been normalized to 1 at $4.12 \mu\text{m}$.

A.3. Results: Comparison of $4\ \mu\text{m}$ Bands in SpeX and NIMS Spectra

In this appendix, we compare continuum-divided $4\ \mu\text{m}$ bands detected in the grand average leading and trailing hemisphere SpeX spectra to average spectra extracted from three different NIMS observations of Callisto (Figure A3).

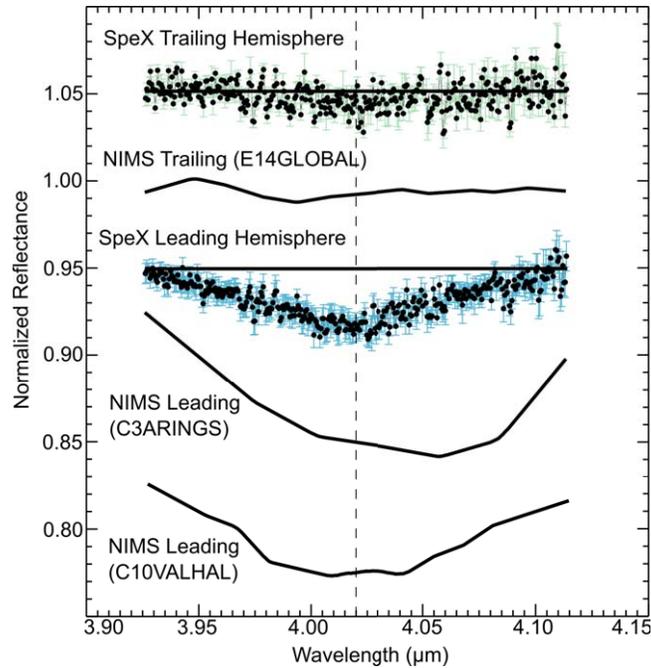


Figure A3. Comparison between continuum-divided $4\ \mu\text{m}$ bands observed in the grand average SpeX spectra of Callisto's leading and trailing hemispheres (blue and green error bars, respectively) to continuum-divided $4\ \mu\text{m}$ bands observed in average NIMS spectra collected over Callisto's trailing hemisphere (E14GLOBAL), proximal to the Asgard impact basin (C3ARINGS), and proximal to the Valhalla impact basin (C10VALHAL), offset vertically for clarity. The dashed line at $4.02\ \mu\text{m}$ marks the central wavelength position of the $4\ \mu\text{m}$ band identified in the SpeX spectra of Callisto.

A.4. Results: Subtle Absorption Bands in the Grand Average IRTF/SpeX Spectra

In this appendix, we present grand average Callisto spectra, zoomed in on the $\sim 2.85\text{--}4.15\ \mu\text{m}$ wavelength region, to highlight some of the subtle absorption features hinted at in these data (Figure A4). The grand average spectra display subtle bands centered near $2.97\ \mu\text{m}$, which appears to be stronger in the mean leading hemisphere spectrum. This band was previously detected by NIMS, and attributed to trapped OH (e.g., Moore et al. 2004). Alternatively, the N–H asymmetric stretch fundamental band (ν_3) is centered near $2.97\ \mu\text{m}$ and could explain the presence of this feature as well. The leading hemisphere spectrum also displays a subtle band near $3.05\ \mu\text{m}$, which was originally detected by prior ground-based observations and attributed to the presence of ammonium-bearing clays (e.g., Calvin & Clark 1993). We also note a weak band near $3.4\ \mu\text{m}$ on Callisto’s leading hemisphere, which was detected previously by ground-based observers and attributed to the presence of small H_2O ice grains (e.g., Calvin & Clark 1993), as well as by NIMS, attributed to the presence of organics (e.g., McCord et al. 1998). This subtle $3.4\ \mu\text{m}$ feature could also result from carbonyl sulfide (OCS;

Coblentz Society 1964). The detection of these possible N-bearing features near 2.97 and $3.05\ \mu\text{m}$, and a possible C- and S-bearing band near $3.4\ \mu\text{m}$, primarily on Callisto’s leading hemisphere, is consistent with the stronger $4.6\ \mu\text{m}$ band on Callisto’s leading side (Figure 1). The $4.6\ \mu\text{m}$ band has been attributed to CN-bearing species by some studies (e.g., McCord et al. 1998), and we suggest that it could also stem from carbon disulfide (CS_2 ; Coblentz Society 1964). Similar to our preferred hypotheses for the origin of the $4\ \mu\text{m}$ band, perhaps the features centered near 2.97 , 3.05 , 3.4 , and $4.6\ \mu\text{m}$ result from constituents delivered in dust grains from the retrograde irregular satellites. Alternatively, perhaps the features near 3.4 and $4.6\ \mu\text{m}$ result from radiolytic products that are formed by magnetospheric charged particles bombarding Callisto’s surface.

We also note the presence of a subtle band centered near $3.75\ \mu\text{m}$, which is slightly stronger on Callisto’s trailing side. This band has not been previously noted on Callisto in either ground-based or NIMS spectra. Although a $3.78\ \mu\text{m}$ band has been detected on Europa (Trumbo et al. 2017), the $3.75\ \mu\text{m}$ band we have detected on Callisto is much narrower and weaker, and it is uncertain whether these two bands might share a common origin.

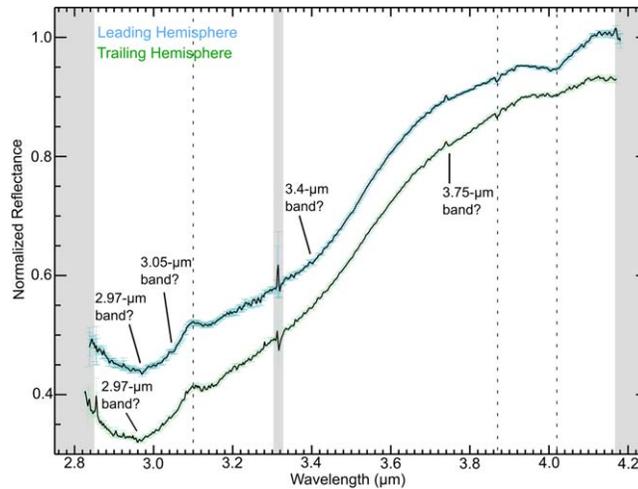


Figure A4. Grand average spectra, and 1σ uncertainties, of Callisto’s leading and trailing hemisphere (blue and green error bars, respectively). These data are normalized to 1 at $4.12\ \mu\text{m}$, and offset vertically for clarity. The dashed lines at 3.1 , 3.87 , and $4.02\ \mu\text{m}$ highlight the central wavelengths of the H_2O ice Fresnel peak and the $3.9\ \mu\text{m}$ and $4\ \mu\text{m}$ bands, respectively. Subtle absorption bands hinted at in the grand average leading hemisphere spectrum data, centered near 2.97 , 3.05 , and $3.4\ \mu\text{m}$, and in the grand average trailing hemisphere spectrum, centered near 2.97 and $3.75\ \mu\text{m}$, are indicated. Both spectra have been lightly smoothed using a 9 pixel-wide boxcar function. Wavelength regions where strong telluric bands are present are shown as gray-toned zones.

Table A1
F-Test Analysis of the Longitudinal Distribution of the 4 μ m Band

Band	Band Measurement	One Tailed <i>F</i> -test Ratio	Sample Size (<i>n</i>)	Mean Model Degree of Freedom (<i>n</i> - 1)	Sinusoidal Model Degree of Freedom (<i>n</i> - 3)	Probability (<i>p</i>)	Reject Null Hypothesis?
4 μ m	Area	34.62	8	7	5	0.0006	Yes
	Depth	25.22	8	7	5	0.0013	Yes

A.5. Results: 4 μ m Band Parameter Measurements

In this appendix, we report the results of our *F*-test analysis (Table A1).

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References

- Bennett, C. J., Pirim, C., & Orlando, T. M. 2013, *ChRv*, **113**, 9086
 Botke, W. F., Vokrouhlický, D., Nesvorný, D., & Moore, J. M. 2013, *Icar*, **223**, 775
 Burns, J. A., Lamy, P. L., & Soter, S. 1979, *Icar*, **40**, 1
 Calvin, W. M., & Clark, R. N. 1993, *Icar*, **104**, 69
 Carlson, R., Anderson, M., Johnson, R., Schulman, M., & Yavrouian, A. 2002, *Icar*, **157**, 456
 Carlson, R., Smythe, W., Baines, K., et al. 1996, *Sci*, **274**, 385
 Cartwright, R. J., Beddingfield, C. B., Nordheim, T. A., et al. 2020a, *ApJL*, **898**, L22
 Cartwright, R. J., Emery, J. P., Grundy, W. M., et al. 2020b, *Icar*, **338**, 113513
 Cartwright, R. J., Emery, J. P., Pinilla-Alonso, N., et al. 2018, *Icar*, **314**, 210
 Cartwright, R. J., Emery, J. P., Rivkin, A. S., Trilling, D. E., & Pinilla-Alonso, N. 2015, *Icar*, **257**, 428
 Coblenz Society 1964, NIST Chemistry WebBook, NIST Standard Reference Database Number 69, <https://webbook.nist.gov/chemistry/>
 Cruikshank, D., Bell, J., Gaffey, M., et al. 1983, *Icar*, **53**, 90
 Cruikshank, D. P. 1980, *Icar*, **41**, 240
 Cushing, M. C., Vacca, W. D., & Rayner, J. T. 2004, *PASP*, **116**, 362
 Emery, J. P., Burr, D. M., & Cruikshank, D. P. 2010, *AJ*, **141**, 25
 Fanale, F. P., Brown, R. H., Cruikshank, D. P., & Clake, R. N. 1979, *Natur*, **280**, 761
 Graykowski, A., & Jewitt, D. 2018, *AJ*, **155**, 184
 Hage, W., Liedl, K. R., Hallbrucker, A., & Mayer, E. 1998, *Sci*, **279**, 1332
 Hand, K. P., Carlson, R. W., & Chyba, C. F. 2007, *AsBio*, **7**, 1006
 Hendrix, A. R., Hurford, T. A., Barge, L. M., et al. 2019, *AsBio*, **19**, 1
 Hendrix, A. R., & Johnson, R. E. 2008, *ApJ*, **687**, 706
 Hibbitts, C., Klemaszewski, J., McCord, T., Hansen, G., & Greeley, R. 2002, *JGRE*, **107**, 5084
 Hibbitts, C., McCord, T., & Hansen, G. 2000, *JGR*, **105**, 22541
 Hudson, R., & Moore, M. 2001, *JGR*, **106**, 33275
 Jiménez-Escobar, A., & Muñoz Caro, G. 2011, *A&A*, **536**, A91
 Lane, A. L., & Chapman, C. R., Bierhaus, E. B., et al. 2004, in *Jupiter: The Planet, Satellites and Magnetosphere*, ed. F. Bagenal, T. E. Dowling, & W. B. McKinnon (Cambridge: Cambridge Univ. Press), 485
 Kliore, A., Anabtawi, A., Herrera, R., et al. 2002, *JGRA*, **107**, 1407
 Lane, A. L., & Domingue, D. L. 1997, *GeoRL*, **24**, 1143
 Liuzzo, L., Simon, S., & Regoli, L. 2019, *P&SS*, **166**, 104726
 Lomax, R. G., & Hahs-Vaughn, D. L. 2013, *An Introduction to Statistical Concepts* (New York: Routledge)
 Mastrapa, R., Sandford, S., Roush, T., Cruikshank, D., & Dalle Ore, C. 2009, *ApJ*, **701**, 1347
 McCord, T. a., Carlson, R., Smythe, W., et al. 1997, *Sci*, **278**, 271
 McCord, T. a., Hansen, G., Clark, R., et al. 1998, *JGR*, **103**, 8603
 Moore, J. M., Chapman, C. R., Bierhaus, E. B., et al. 2004, in *Jupiter: The Planet, Satellites and Magnetosphere*, ed. F. Bagenal, T. E. Dowling, & W. B. McKinnon (Cambridge: Cambridge Univ. Press), 397
 Moore, M., Hudson, R., & Carlson, R. 2002, *BAAS*, **34**, 902
 Moore, M., Hudson, R., & Carlson, R. 2007, *Icar*, **189**, 409
 Morbidelli, A., & Nesvorný, D. 2020, in *The Trans-Neptunian Solar System*, ed. D. Pralnik, M. A. Barucci, & L. Young (Amsterdam: Elsevier), 25
 Morrison, D., Morrison, N. D., & Lazarewicz, A. R. 1974, *Icar*, **23**, 399
 Nash, D. B., & Betts, B. H. 1995, *Icar*, **117**, 402
 Noll, K. S., Johnson, R. E., McGrath, M. A., & Caldwell, J. J. 1997, *GeoRL*, **24**, 1139
 Nyquist, R. A., Leugers, M. A., & Putzig, C. L. 1997, *The Handbook of Infrared and Raman Spectra of Inorganic Compounds and Organic Salts* (New York: Academic)
 Pollack, J. B., Witteborn, F. C., Erickson, E. F., et al. 1978, *Icar*, **36**, 271
 Rayner, J., Toomey, D., Onaka, P., et al. 2003, *PASP*, **115**, 362
 Salama, F., Allamandola, L., Witteborn, F., et al. 1990, *Icar*, **83**, 66
 Schneider, N. M., & Bagenal, F. 2007, *Io After Galileo* (Berlin: Springer)
 Spencer, J. R. 1987, PhD thesis, Univ. of Arizona, <https://repository.arizona.edu/handle/10150/184098>
 Szabó, G. M., Ivezić, Ž., Jurić, M., & Lupton, R. 2007, *MNRAS*, **377**, 1393
 Tamayo, D., Burns, J. A., & Hamilton, D. P. 2013, *Icar*, **226**, 655
 Trumbo, S. K., Brown, M. E., Fischer, P. D., & Hand, K. P. 2017, *AJ*, **153**, 250
 Wong, I., & Brown, M. E. 2016, *AJ*, **152**, 90